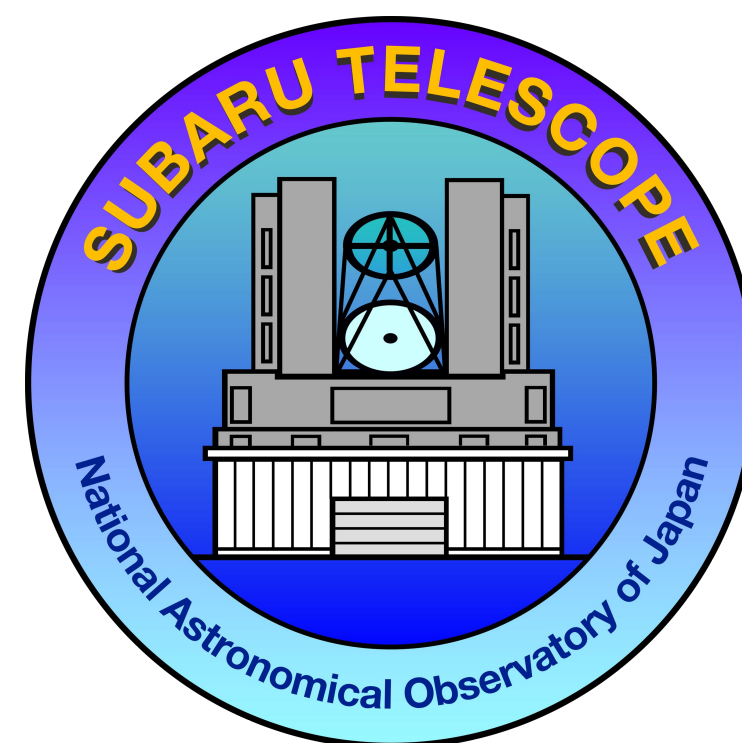




Diffraction-Limited Visible-Light Imaging with VAMPIRES

Miles Lucas^a, Barnaby Norris^e, Olivier Guyon^{b,c,d}, Michael Bottom^a, Peter Tuthill^e, Kyohoon Ahn^b, Vincent Deo^b, Julien Lozi^b, Sebastien Vievard^b
^a University of Hawai'i, ^b Subaru telescope, ^c Astrobiology Center, ^d University of Arizona, ^e University of Sydney



PI-type instrument, AVAILABLE FOR OPEN USE

Contact: mdlucas@hawaii.edu

OVERVIEW

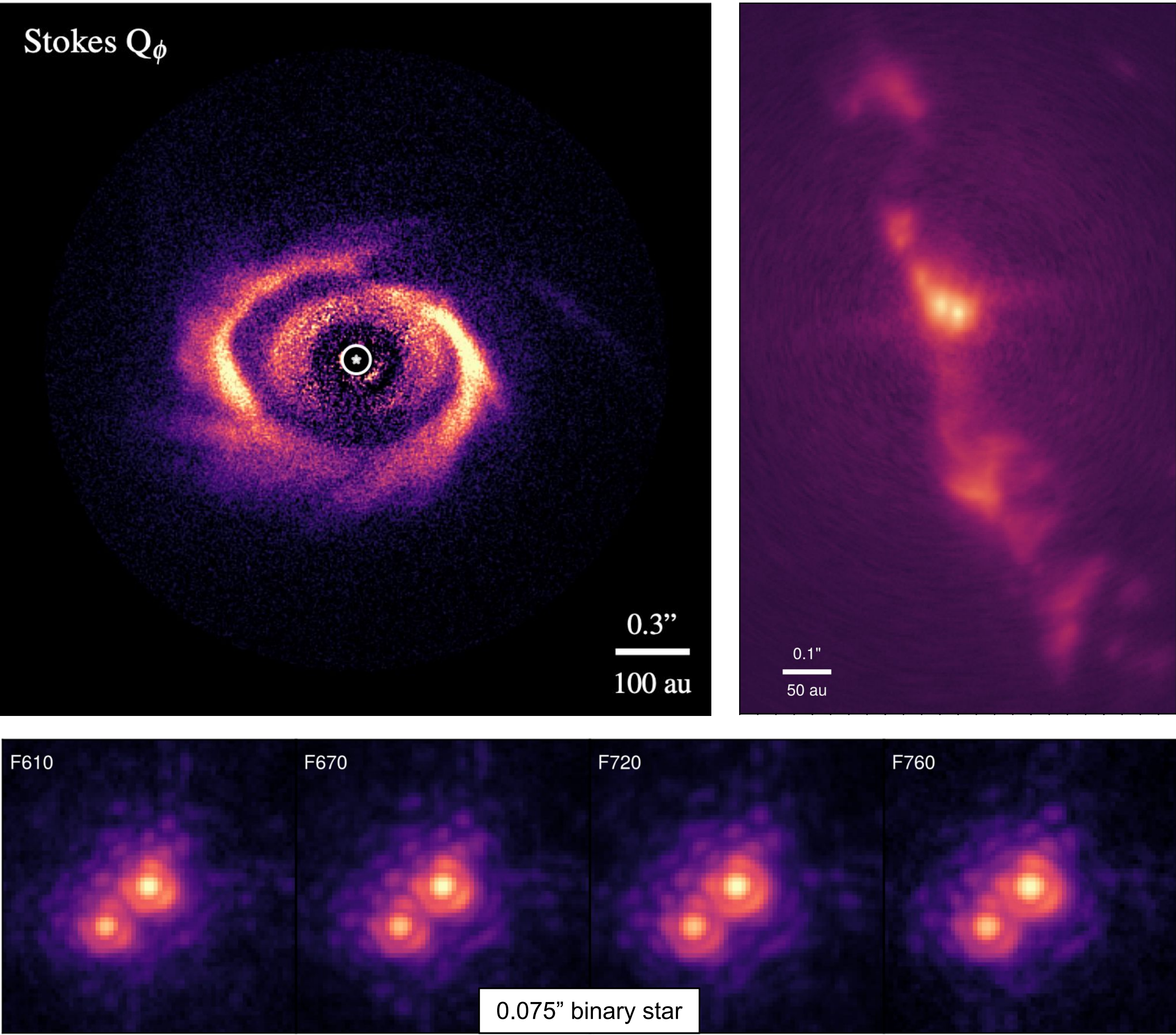
VAMPIRES is SCEXAO's **visible light (600 nm - 800 nm) dual-beam imager**, optimized for high angular resolution and high contrast imaging. VAMPIRES includes polarimetric, multi-band, sparse aperture masking and coronagraphic capabilities.

VAMPIRES benefits from the two-stage AO correction (AO188/AO3k + SCEXAO), and operates at the diffraction limit, providing ~0.02" angular resolution.

PI: Peter Tuthill, University of Sydney, Australia.
Instrument team: University of Hawaii, University of Sydney and Subaru Telescope.

VAMPIRES operates simultaneously with near-IR SCEXAO-fed instrumentation, including CHARIS.

SCIENCE GOALS AND CAPABILITIES



Instrument modes and capabilities:

- 3" x 3" FOV, 6 mas/px plate scale
- 5 Selectable 50 nm broadband filters plus multiband-imaging mode
- High angular resolution with AO188 + SCEXAO wavefront control and correction
- Optimized for polarimetry; can use HWP alone or HWP + FLC for fast switching
- Coronagraphy (4 classic Lyot masks + dgVVC)
- Sparse aperture masking
- Narrowband spectral differential imaging (Halpha, SII)

Simple VAMPIRES results (left):

- Polarimetric scattered-light imaging of circumstellar disks (top left; HD 34700 disk polarized intensity at 750 nm)
- Halpha/SII narrowband imaging (top right; R Aqr emission nebula in Halpha)
- Halpha spectral differential imaging for accreting protoplanets (Uyama+2020)
- Sub-diffraction-limited imaging of mass-loss shells around giant stars using non-redundant aperture mask interferometry (Norris+2015)
- High-angular resolution imaging of stellar companions (bottom; 75 mas binary HD 204827 in multiband imaging mode)

RECENT HIGHLIGHTS

Major recent hardware upgrades :

- EMCCD sensors were replaced with ultra-low read noise CMOS detectors (photon number resolving) **improved dynamic range and speed.**
- New **multiband imaging mode** for simultaneous imaging through multiple broadband filters, compatible with polarization differential imaging and coronagraphy.
- New achromatic FLC (improved polarimetric efficiency at the shortest and longest wavelengths)
- Deployed double-grating vector vortex coronagraph (dgVVC; PI: D. Doelman) for improved contrast at small angular separations
- Deployed redundant apodizing pupil mask (RAP; PI: L. Leboulex) for increased resilience to wavefront errors
- Data processing pipeline available

FUTURE PLANS

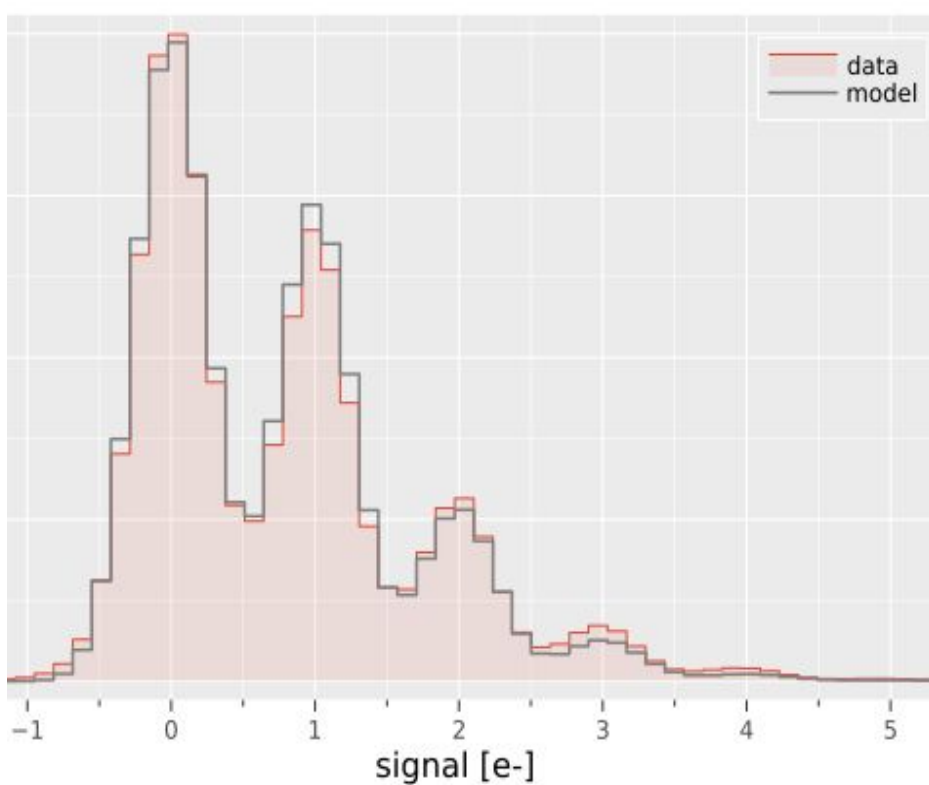
- Closed-loop pointing control of dgVVC (QACITS) + on-sky testing
- On-sky testing of RAP
- Development of coronagraphic low-order wavefront sensing loop
- Focal plane wavefront sensing development (fast and furious, LAPD, LDFC, petalometry)
- Halpha velocity mapping with narrowband dichroic image slicer

TECHNICAL OVERVIEW

Detectors

Two identical CMOS detectors, (*Hamamatsu ORCA-Quest*)

- 0.22 e- RMS read noise @ 40+ Hz
- 0.4 e- RMS read noise @ 500+ Hz
- 7.2 us to 1800 s DIT
- Hardware sync between both cameras and FLC



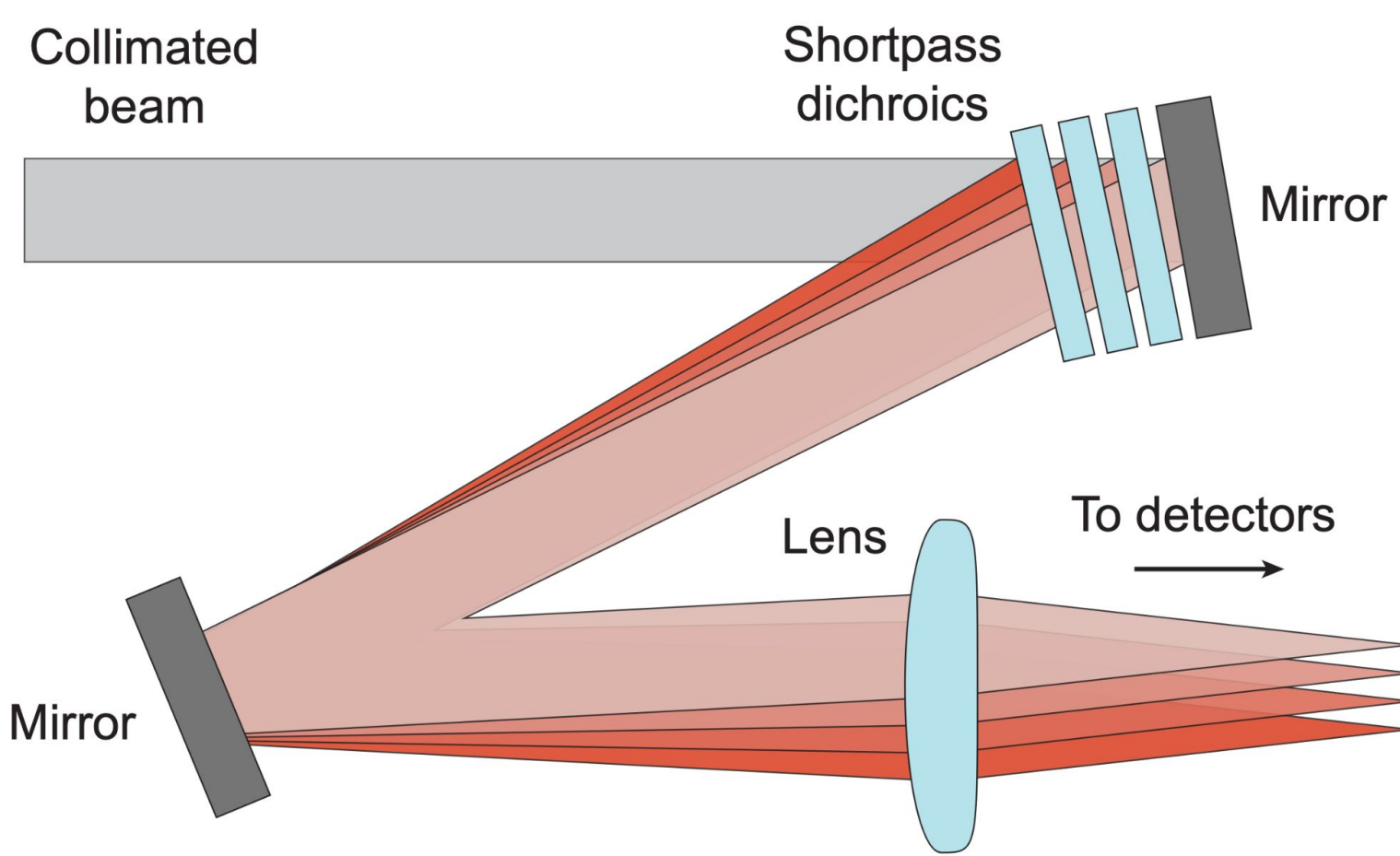
Control System

- Asynchronous python control code
- PyGame camera viewers
- Real time control with CACAO and MILK

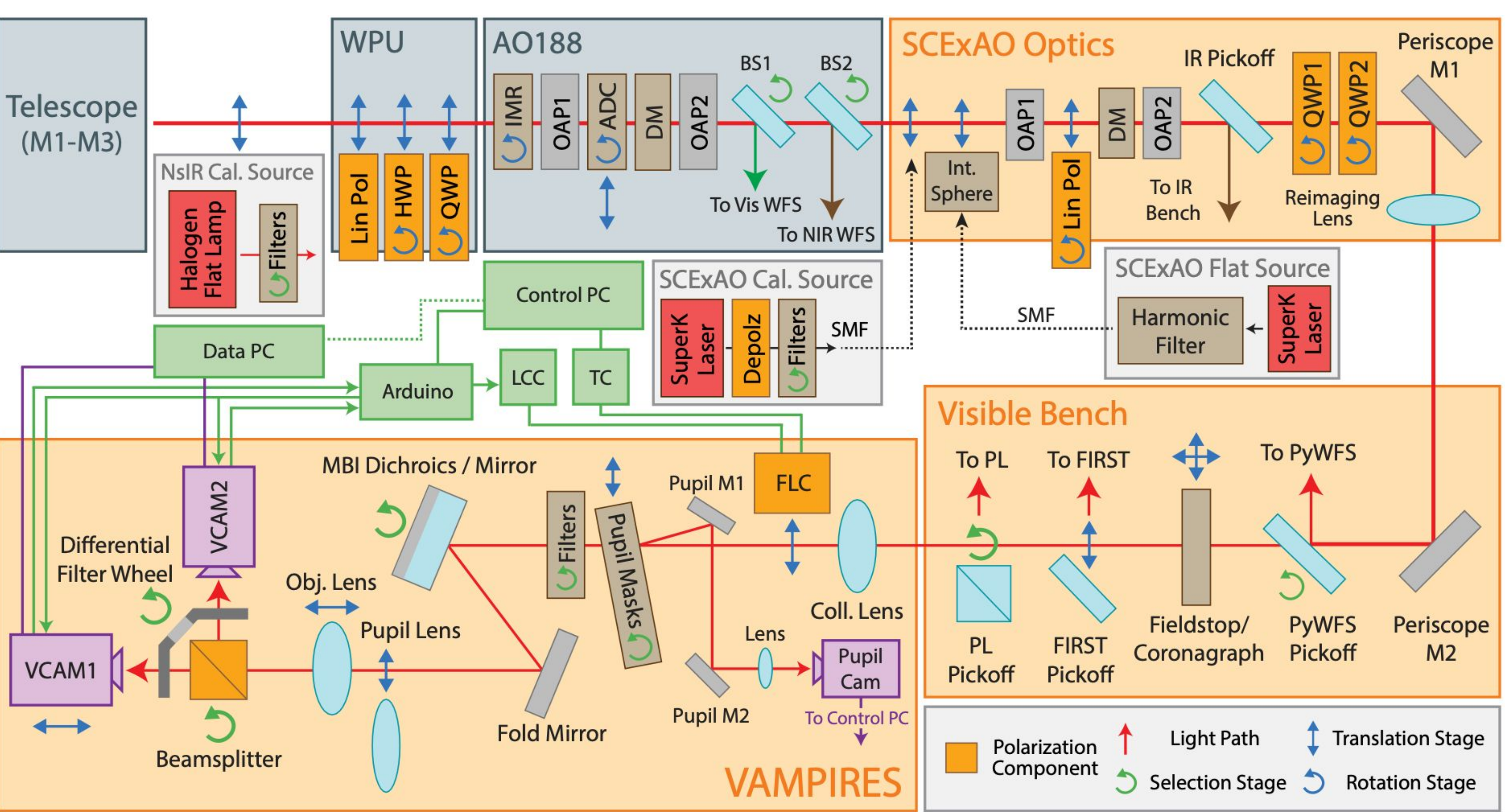
Precision timing: VAMPIRES science frames can be correlated with WFS telemetry and other high-speed cameras.

Multiband Imaging

- 4 fields imaged simultaneously
- Stack of low AOI dichroics



INSTRUMENT SCHEMATIC



Acknowledgements

The authors wish to recognize and acknowledge the very significant cultural role and reverence that the summit of Maunakea has always had within the Hawaiian community. We are most fortunate to have the opportunity to conduct observations from this mountain.

The development of SCEXAO was supported by the National Astronomical Observatory of Japan (NAOJ), the Astrobiology Center of the National Institutes of Natural Sciences, Japan, the Subaru Telescope, the Japan Society for the Promotion of Science (Grant-in-Aid for Research #23340051, #26220704, #23103002, #19H00703 & #19H00695), and the Mt Cuba Foundation.

Wavefront control R&D activities received support from the Heising-Simons foundation and NASA (Grant #80NSSC19K0336).